

General Features of the Linear Crystalline Morphology of Accretion Disks

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Abstract In this paper, we analyze the so-called Master Equation of the linear backreaction of a plasma disk in the central object magnetic field, when small scale ripples are considered. This study allows to single out two relevant physical properties of the linear disk backreaction: i. the appearance of a vertical growth of the magnetic flux perturbations; ii. the emergence of sequence of magnetic field O -points, crucial for the triggering of local plasma instabilities. We first analyze a general Fourier approach to the solution of the addressed linear partial differential problem. This technique allows to show how the vertical gradient of the backreaction is, in general, inverted with respect to the background one. Instead, the fundamental harmonic solution constitutes a specific exception for which the background and the perturbed profiles are both decaying. Then, we study the linear partial differential system from the point of view of a general variable separation method. The obtained profile describes the crystalline behavior of the disk. Using a simple rescaling, the governing equation is reduced to the second order differential Whittaker equation. The zeros of the radial magnetic field are found by using the solution written in terms Kummer functions. The possible implications of the obtained morphology of the disk magnetic profile are then discussed in view of the jet formation.

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1 Introduction

The Standard Model of accretion in astrophysics [?,?] is based on the idea that the angular momentum transport across the plasma disk is ensured by effective viscosity, due to the turbulence phenomena triggered by the magneto-rotational instability (MRI) [?], see also Refs.[?,?,?]. However, the necessity to account for a magnetic field, due to the central object, together with the long-wavelength response of the plasma disk, implies that the azimuthal component of the generalized Ohm law is consistent only postulating an anomalous plasma resistivity, see Refs.[?,?]. This effective resistivity makes the disk magnetic field limited in amplitude, because of its diffusive nature [?]. As a result, the possibility to deal with the high magnetic field strength required for the jet generation leads to investigate specific scenarios [?].

In Refs.[?,?] (see also Ref.[?]), it has been investigated the possibility for a small scale backreaction of the plasma disk to the magnetic field of the central object, outlining the emergence of a crystalline morphology of the magnetic field microscales, *i.e.*, a local radial oscillation of the magnetic flux function. In particular, in the limit of a linear backreaction, in Ref.[?] it has been written down a Mas-

ter Equation for the magnetic flux function, which is the starting point of the present investigation. The relevance of this new paradigm relies on the possibility to use the O -points of the resulting configuration as the sites where the jet formation could start, as discussed in Refs.[?,?].

Here, we develop a detailed general analysis of the linear Master Equation for the crystalline profile of the plasma disk, outlining two main significant features: *i*) performing a general Fourier analysis, we clarify how the decreasing behavior of the plasma backreaction with the vertical quote (discussed in Refs.[?,?], see also Ref.[?]) is actually a peculiar property of the fundamental harmonics in the radial oscillation; *ii*) by analyzing the Master Equation in the framework of a separable solution, we clarify the existence of regions of the parameter space, where a relevant sequence of O -points of the magnetic configuration can take place.

Both the results mentioned above have a precise physical implication. The point *i*) shows how the higher order harmonics of the radial oscillation of the flux function are actually associated to an inverse behavior of the vertical gradient between the background and the linear perturbation. This issue calls attention for its possible implications on the morphology of the non-linear backreaction which is

associated with the fragmentation of the disk in a ring-like morphology, see Ref.[?] and also Refs.[?,?,?]. The point *ii)* demonstrates that the topology of the induced radial component of the magnetic field can be characterized by a sequence of *O*-points, where such a radial component vanishes (noting that the dipole-like magnetic field of the central object is essentially vertical on the equatorial plane). This feature, absent in the analysis in Refs.[?,?], suggests that the magnetic microstructure, emerging when the perturbation scale is sufficiently small (*i.e.*, for sufficiently high values of the plasma β parameter), could be an interesting mechanism in triggering sites for the jet formation. It is worth noting that we also outline the existence of a parameter region for the model in which both the radial and the vertical dependence of the backreaction magnetic surfaces oscillate.

2 Standard model for accretion

The basic idea for accretion onto a compact object has been formulated in Ref.[?] and Ref.[?], where essentially thin disk configurations were considered. Actually, thin accretion disks are commonly present around stellar and black hole systems and they have the advantage to elucidate the accretion features in a one-dimensional model, *i.e.*, essentially the radial dependence on the r coordinate is described, being the vertical one (z) averaged out of the problem and the azimuthal dependence on ϕ is forbidden by the axial symmetry of the steady disk configuration [?].

For such a steady axially symmetric one dimensional representation of a thin disk configuration, accreting material onto a compact object, having a mass M and an almost vertical magnetic field $\mathbf{B} = B\hat{\mathbf{e}}_z$, is determined via the implementation of the mass conservation equation and of the momentum conservation ones.

If we define the disk accretion rate as

$$\dot{M} \equiv -2\pi r \Sigma u_r, \quad (1)$$

where Σ and u_r denote the superficial density and the radial infalling velocity of the fluid, respectively, then the mass conservation equation reduces to the following condition

$$d\dot{M}/dr = 0. \quad (2)$$

The radial component of the momentum conservation equation splits into two components, one stating that the disk angular velocity ω is essentially equal to the Keplerian one ω_K and the advection and the pressure gradient balance each other separately, *i.e.*,

$$\omega(r) \simeq \omega_K = \sqrt{GM/r^3}, \quad (3)$$

$$\rho_0 u_r \frac{du_r}{dr} + \frac{dp_0}{dr} = 0, \quad (4)$$

$\rho_0(r)$ and $p_0(r)$ denoting the values of the mass density and of the pressure, taken on the equatorial plane of the disk.

The vertical component of the momentum equation determines the decay law of the mass density far away the

equatorial plane. The details of such decay depend on the particular equation of state we adopt for the fluid, but a reliable expression, valid both for the general polytropic and isothermal case, reads as

$$\rho(r, z) \simeq \rho_0(r) \left(1 - \frac{z^2}{H^2}\right), \quad (5)$$

where $H(r)$ denotes the half-width of the disk and if it is thin we must have $H/r \ll 1$. The spatial scale of the disk vertical extension is fixed in the vertical gravostatic equilibrium via the ratio between the sound velocity to the Keplerian angular velocity, *i.e.*, $H \sim v_s/\omega_K$. The sound velocity can be reliably estimated via the relation $v_s \sim \sqrt{K_B T/m}$, being T the disk temperature, K_B the Boltzmann constant and eventually m the proton mass.

Finally, the azimuthal component of the momentum conservation equation, which regulates the angular momentum transport across the disk, provides the following expression of the accretion rate:

$$\dot{M} = 6\pi\eta_v H, \quad (6)$$

here η_v denotes the shear viscosity coefficient, associated to the friction of the disk layers in differential rotation with ω_K . This viscosity effect can not be due to the kinetic conditions of the plasma, which is actually quasi-ideal (see Ref.[?]), and in Ref.[?] η_v has been justified via a turbulent behavior emerging in the disk spatial microscales. Such a coefficient can be estimated by comparing the exact expression of the (r, ϕ) component of the viscous stress tensor τ_{ij} with its interpretation in terms of the correlation function of the turbulent radial and azimuthal velocity component v_r and v_ϕ , respectively, *i.e.*, we have

$$\tau_{r\phi} = \eta_v r \frac{d\omega_K}{dr} = -\frac{3}{2}\eta_v \omega_K = -\langle \rho_0 v_r v_\phi \rangle = -\alpha \rho_0 v_s^2, \quad (7)$$

from which we get the basic Shakura expression

$$\eta_v = \frac{2}{3}\alpha \rho_0 v_s H. \quad (8)$$

The parameter $\alpha \leq 1$ has been introduced to phenomenologically account that all the supersonic velocity fluctuations are unavoidably damped in the turbulent regime.

By combining this expression for η_v with the relation (6), we arrive to the final expression for the accretion rate of the disk in terms of basic quantities (particle density, temperature, central body mass), namely

$$\dot{M} = 4\pi\alpha\rho_0 v_s H^2 = 4\pi\alpha\rho_0 \omega_K H^3. \quad (9)$$

The possibility to deal with a turbulent behavior of the plasma requires that a sufficiently significant instability can affect the equilibrium configuration. Since in the case of a Keplerian disk the steady profile turns out to be stable [?], after some years of investigation, the source of turbulence has been identified in the MRI [?,?]. Firstly derived in Ref.[?] and Ref.[?], MRI is an Alfvénic instability taking place when the magnetic field intensity within

the plasma is below a given threshold. In particular, for a Keplerian disk, the simplest condition to deal with MRI reads as follows

$$\omega_A \equiv kv_A < \sqrt{3} \omega_K, \quad (10)$$

v_A being the Alfvén speed and k the perturbation wavenumber. In correspondence to an assigned value of the magnetic field, a minimal spatial scale exists $\lambda = 2\pi/k$ for which MRI holds. Such a scale, say λ_{min} is easily identified as

$$\lambda_{min} \simeq \frac{2\pi}{\sqrt{3}} \frac{v_A}{v_s} H \sim \frac{H}{\sqrt{\beta}}, \quad (11)$$

β denoting the standard plasma parameter, which in astrophysics can take also rather large values.

In Ref.[?], it has been shown that just on the small scale λ_{min} , there stated as $\sim v_A/\omega_K$, the plasma backreaction to the central body can also take a very different nature with respect turbulence. In fact, a steady perturbation to the equilibrium is allowed for which the correction of the centripetal force associated to the differential rotation, is directly linked to the perturbed Lorentz force raised in the plasma by the emergence of current densities on very small scales. We observe that the background magnetic field is current-free since it is due to the central body magnetosphere and, within the disk, it is a vacuum field. In Ref.[?], this idea of a microscopic crystalline morphology of the perturbed magnetic field, *i.e.*, its radial oscillating behavior on the scale λ_{min} , is implemented also in the limit of a non-linear backreaction. As a result, it is possible to show that the disk morphology is decomposed in a series of microscopic ring-like structures, *i.e.*, the mass density of the disk acquires periodic nodes, see also the global model developed in Ref.[?].

In the simplest case of a linear small scale backreaction within the disk, the magnetic flux function is corrugated and this perturbed profile is associated to a linear two-dimensional partial differential equation, dubbed in what follows as Master Equation for the disk crystalline structure. We will discuss in detail the morphology of the solutions associated to this equation in order to extract information about the small scale properties of the disk steady profile on which MRI can develop. We conclude by observing that in Ref.[?], it was argued how the presence of a small scale crystalline disk backreaction and the associated relevant micro-current density can have significant implications on the problem of the “anomalous” disk resistivity, required to account for the observed accretion rates in systems like X-ray binary stars.

3 Physical model

We consider a steady and axialsymmetric thin disk configuration, characterized by a magnetic field \mathbf{B} having the following poloidal form

$$\mathbf{B} = -r^{-1} \partial_z \psi \hat{\mathbf{e}}_r + r^{-1} \partial_r \psi \hat{\mathbf{e}}_z, \quad (12)$$

where we use the aforementioned standard cylindrical coordinates (r, ϕ, z) ($\hat{\mathbf{e}}_r$, $\hat{\mathbf{e}}_\phi$ and $\hat{\mathbf{e}}_z$ being their versors) and

$\psi(r, z)$ is the magnetic flux function. The disk also possesses a purely azimuthal velocity field \mathbf{v} , *i.e.*,

$$\mathbf{v} = \omega(\psi) r \hat{\mathbf{e}}_\phi, \quad (13)$$

where the angular velocity ω is a function of ψ because we are in the range of validity of the so-called co-rotation theorem [?], *i.e.*, Eqs.(12) and (13) holds together with the stationary and axial symmetry hypotheses.

Since, we are assuming the plasma disk is quasi-ideal (actually it is true in many range of observed mass density and temperature), we neglect the poloidal velocities, especially the radial component, which are due to effective dissipation, according to the Shakura idea of accretion discussed in the previous Section. Furthermore, we observe that the co-rotation condition (13) prevents the emergence of an azimuthal magnetic field component via the dynamo effect.

We now split the magnetic flux function around a fiducary radius $r = r_0$, as follows

$$\psi = \psi_0(r_0) + \psi_1(r_0, r - r_0, z), \quad (14)$$

where ψ_0 is the vacuum contribution of the central object around which the disk develops (essentially a vertical magnetic field comes out from the dipole-like nature of the field and from the thinness of the disk), while ψ_1 is a small (still steady) correction, here considered of very small scale with respect to the background quantities. By other words, we are studying a small backreaction of the plasma which is embedded in the central object magnetic field, whose spatial (radial and vertical) scales are sufficiently small to produce non-negligible local currents in the disk.

According to Eqs.(14) and (12), also the magnetic field is expressed as $\mathbf{B} = \mathbf{B}_0 + \mathbf{B}_1$. According to the validity of the co-rotation theorem, at any order of perturbation of the steady configuration, we expand the angular velocity as follows

$$\omega(\psi) \simeq \omega_0(\psi_0) + (d\omega/d\psi)_{\psi=\psi_0} \psi_1. \quad (15)$$

In Ref.[?] (see also Refs.[?,?]), it was shown that, in the linear regime, *i.e.*, $|\mathbf{B}_1| \ll |\mathbf{B}_0|$, the equilibrium configuration near r_0 reduces to the radial equilibrium only, which, at the zeroth and first order in ψ , gives the following two equations

$$\omega_0(\psi_0) = \omega_K, \quad (16)$$

$$\partial_r^2 \psi_1 + \partial_z^2 \psi_1 = -k_0^2 (1 - z^2/H^2) \psi_1, \quad (17)$$

respectively. We recall that H denotes the half-depth of the disk and $k_0 \sim 1/\lambda_{min}$ is the typical wavenumber of the small scale backreaction, taking the explicit form

$$k_0^2 \equiv 3\omega_K^2/v_A^2, \quad (18)$$

and, for a thin isothermal disk, it results $k_0 H = \sqrt{3}\beta \equiv 1/\epsilon$. As postulated above, in order to deal with small scale perturbations, we have to require that the value of β is sufficiently large.

In order to study the solutions of Eq.(17), it is convenient to introduce dimensionless quantities, as follows

$$Y \equiv \frac{k_0 \psi_1}{\partial_{r_0} \psi_0}, \quad x \equiv k_0 (r - r_0), \quad u = z/\delta, \quad (19)$$

where $\delta^2 = H/k_0$. Hence, we get

$$\partial_x^2 Y + \epsilon \partial_u^2 Y = -(1 - \epsilon u^2) Y, \quad (20)$$

which, in what follows we dub the Master Equation for the crystalline structure of the plasma disk. We recall that, according to Eq.(5), the u variable takes a finite range across the thin disk configuration. Moreover, it is also easy to check the validity of the relations

$$B_z = B_{0z} (1 + \partial_x Y), \quad (21)$$

$$B_r \equiv B_{1r} = -B_{0z} \sqrt{\epsilon} \partial_u Y, \quad (22)$$

where the existence of linear perturbation regime requires $|Y| \ll 1$.

It is relevant to investigate the solutions of Eq.(20) in view of determining the physical conditions under which the crystalline structure, discussed in Refs.[?,?], can actually take place.

4 A viable solution to the Master Equation

We investigate the solution of the Eq.(20) for the crystalline morphology where we consider the regime $\epsilon \ll 1$, corresponding to large values of the β plasma parameter. Since the crystalline structure is associated to a periodicity in the x -dependence, we naturally search a solution in the form

$$Y(x, u) = \sum_{n=1}^{n=\infty} F_n(u) \sin(nx + \varphi_n(u)), \quad (23)$$

where F_n and φ_n are amplitude and phase, respectively. Once set to zero the coefficients of the different trigonometric functions, this expression provides the two following equations

$$2 \frac{dF_n}{du} \frac{d\varphi_n}{du} + F_n \frac{d^2 \varphi_n}{du^2} = 0, \quad (24)$$

$$\epsilon \frac{d^2 F_n}{du^2} - \epsilon F_n \left(\frac{d\varphi_n}{du} \right)^2 = [(n^2 - 1) + \epsilon u^2] F_n. \quad (25)$$

Eq.(24) admits the solution

$$d\varphi_n/du = \Theta/F_n^2, \quad (26)$$

with $\Theta = const.$, which reduces Eq.(25) to the following closed form in F_n

$$\epsilon \frac{d^2 F_n}{du^2} - \epsilon \frac{\Theta^2}{F_n^3} = [(n^2 - 1) + \epsilon u^2] F_n. \quad (27)$$

We observe that, if the periodicity is not associated to the fundamental wavenumber k_0 but to a close one $k = k_0 \delta$

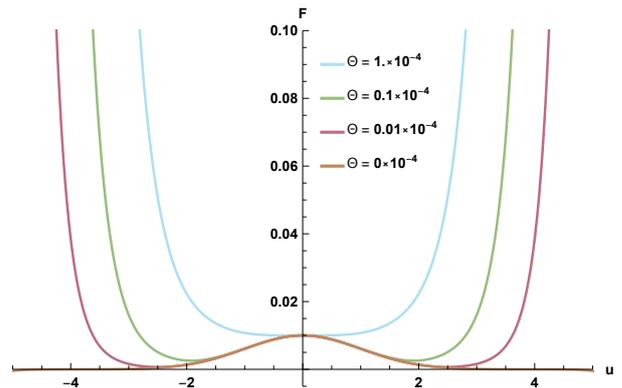


Figure 1. (Color online) Solution of Eq.(28) for $\bar{F} = 0.01$ varying the parameter Θ as indicated in the plot.

(where $0 < \delta \ll 1$), the equation above is simply mapped by replacing the integer n by $n\delta$. For $\epsilon \ll 1$, it is clear that we must have $n = 1$ and then we set $n^2 \delta^2 = 1 - \epsilon$. By this choice, Eq.(27) becomes independent of ϵ , reading as (noting $F \equiv F_1$)

$$\frac{d^2 F}{du^2} - \frac{\Theta^2}{F^3} = -(1 - u^2) F. \quad (28)$$

Let us consider the problem of solving Eq.(28) with boundary condition $F(0) = \bar{F}$, being \bar{F} a generic constant (we consider $0 < \bar{F} \ll 1$ since such an equation is invariant for $F \rightarrow -F$ and because the linear regime must be preserved) and $F'(0) = 0$ (here and in the following the prime denotes u derivatives). This last condition has been selected in order to guarantee the symmetry of the solution with respect to the equatorial plane. The plots of the numerical solution, in correspondence of different values of Θ at fixed \bar{F} , are in Fig.1. Such a solution is a Gaussian for $\Theta = 0$ as derived in Ref.[?], while for $\Theta \neq 0$ it starts to increase for increasing $|u|$ values. It is worth recalling that, when F is greater than unity, the linear approximation fails and the Master Equation is no longer predictive.

We note that, only for $\Theta = 0$, an exponentially decaying behavior of the linear perturbation over the equatorial plane is obtained, just like the background profile [?]. Such a natural similarity in the vertical behavior of the background and of the perturbations was also at the ground of the analyses in Refs.[?,?,?], but the solution derived above demonstrates how this feature is a peculiar property of the simplest case only, *i.e.*, when the phases φ_n are constant quantities.

Clearly the diverging character of the perturbations over the equatorial plane is not really unphysical, both because the solution is calculated in a local region (around fiducial x and u values) and also because the linear theory holds only for small perturbations. Nonetheless, such an increasing behavior is, to some extent, unexpected, because it reverses the natural vertical background profile. This can be physically explained because, while the background profile is fixed by a gravostatic equilibrium, the perturbation morphology has nothing to do with gravity

and it is mainly produced by the balance of the centripetal (say centrifugal) and the Lorentz forces.

However, it is important to stress that for sufficiently small values of Θ a peak of the function F around the equatorial plane is still present. This feature takes place because for small values of Θ the effect of the non-linear term Θ/F^3 is minimized.

From a physical point of view, a process able to induce a crystalline structure in the plasma disk (for instance a sound or a gravitational wave), *i.e.*, a boundary condition able to select the solution associated to a radial corrugation of the magnetic flux function, should be vertically coherent in phase to produce the same vertical gradient of the background profile in the perturbation too. On the contrary, a vertical shear also in the phase of the underlying process is responsible for an inversion of the vertical gradient profile between the background and the perturbation.

The most important physical conclusion of the present analysis is to demonstrate that the crystalline structure is a very general feature of the linearized perturbed equilibrium, but, being associated to small values of ϵ (*i.e.*, large values of the plasma β parameter), the radial oscillation can be associated to a basic wavenumber only, very close to the natural one k_0 .

5 Separable solution of the Master Equation

Since the approach in the previous section gave us the solution $Y = F(u) \sin(x)$ as the only viable case for $\epsilon \ll 1$, it is natural to investigate solutions relying on the separation of variable method. In particular, we will be able to explore the situation in which the wavenumber of the radial oscillation (x -dependence) is a generic constant times the fundamental wavenumber k_0 .

In order to investigate solutions relying on the separation of variable method, the Master Equation (Eq.(20)) can be rewritten using $Y(x, u) = Y_1(x)Y_2(u)$ in the following decoupled form:

$$(d^2Y_1/dx^2)/Y_1 = -\gamma, \quad (29)$$

$$\epsilon(d^2Y_2/du^2)/Y_2 + (1 - \epsilon u^2) = \gamma, \quad (30)$$

where we have introduced the arbitrary constant γ . The first equation has two possible kind of solutions, namely

$$Y_1(x) = C_1 \sin(\sqrt{|\gamma|x}) + D_1 \cos(\sqrt{|\gamma|x}), \quad \gamma > 0, \quad (31)$$

$$Y_1(x) = C_1 \sinh(\sqrt{|\gamma|x}) + D_1 \cosh(\sqrt{|\gamma|x}), \quad \gamma < 0. \quad (32)$$

Eq.(30) can be rewritten in the form of the the Kummer equation of parabolic cylinder functions using the scaling $u \rightarrow u/\sqrt{2}$. In particular, we get

$$\frac{d^2Y_2}{du^2} - \left(\frac{1}{4}u^2 + A\right)Y_2 = 0, \quad (33)$$

where $A = (\gamma - 1)/(2\epsilon)$. Applying, in view of the linearity of (20), the superposition principle, the general solution

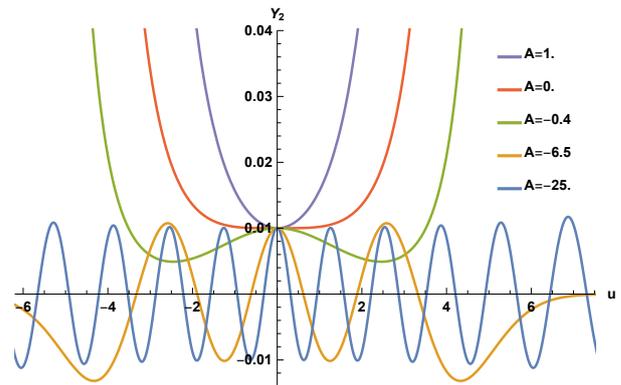


Figure 2. (Color online) Solution of Eq.(33) using $Y_2(0) = 0.01$ and $Y_2'(0) = 0$, by varying the parameter A in $[1, 0, -0.4, -6.5, -25]$ as indicated in the plot.

of the Master Equation is then

$$\begin{aligned} Y(x, u) &= \int d\gamma \chi(\gamma) Y_1(x, \gamma) Y_2(u, \gamma) = \\ &= \int_0^{+\infty} d\gamma \chi(\gamma) (C_1 \sin(\sqrt{|\gamma|x}) + \\ &\quad + D_1 \cos(\sqrt{|\gamma|x}) Y_2(u, \gamma)) + \\ &\quad + \int_{-\infty}^0 d\gamma \chi(\gamma) (C_1 \sinh(\sqrt{|\gamma|x}) + \\ &\quad + D_1 \cosh(\sqrt{|\gamma|x}) Y_2(u, \gamma)), \quad (34) \end{aligned}$$

where $\chi(\gamma)$ is any integrable function.

Let us now analyze the properties of the solution by numerically integrating Eq.(33) using initial conditions $Y_2(0) = 0.01$ and $Y_2'(0) = 0$. We also fix the parameter $\epsilon = 0.01$ and run the γ parameter in order to investigate the morphology of the solution for different values of the constant A . From the numerical analysis, a range of the parameter A emerges (dependent on the choice of ϵ) for which the behavior of the function Y_2 becomes to oscillate, as well sketched in Fig.2. Since these values of A still corresponds to positive values of the constant γ (for our setup, $\gamma \lesssim 0.9$), we obtain a range of the solution of the Master Equation where the radial and vertical dependence of the perturbed magnetic flux function both have an oscillating profile as represented in Fig.3.

Clearly, when γ takes negative values, the x dependence of the function Y is no longer oscillating (while the u -dependence still remains oscillating) and the crystalline morphology of the disk is suppressed. As we shall discuss below, the oscillating nature of the vertical dependence of the perturbed magnetic flux function has the relevant physical implication that, according to the expression for the magnetic field in Eqs.(21) and (22), a series of O -points appears in the disk configuration, in which the radial component B_r vanishes.

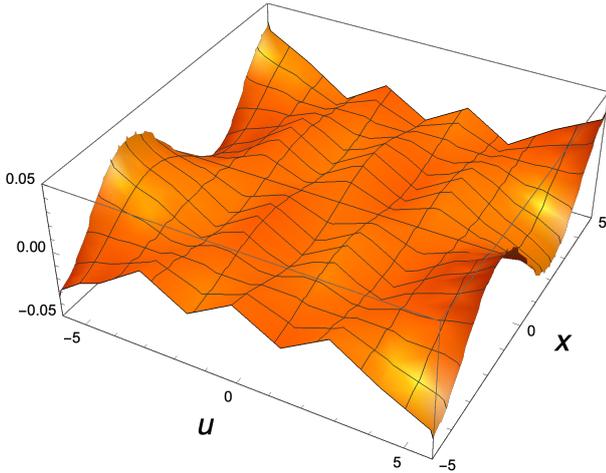


Figure 3. (Color online) Plot of the solution $Y(x, u) = Y_1(x)Y_2(u)$ provided by the system of Eqs.(31) (with $C_1 = 1$ and $D_1 = 0$) and (33) for $Y_2(0) = 0.01$, $Y_2'(0) = 0$, $A = -11$ (i.e., $\gamma = 0.78$).

6 Whittaker equation and O-points

The many different behaviors of the solution of Eq.(33) can be better understood by investigating peculiar properties of the corresponding analytic solutions. In particular, using the following change of variables: $y = u^2/2$ and $Y_2 = (2y)^{-1/4}W(y)$, Eq.(33) rewrites now as

$$\frac{d^2W}{dy^2} + \left(-\frac{1}{4} - \frac{A}{2y} + \frac{3}{16y^2} \right)W = 0, \quad (35)$$

which is a particular case of Whittaker's equation, i.e.,

$$\frac{d^2W}{dy^2} + \left(-\frac{1}{4} + \frac{K}{y} + \frac{1/4 - \mu^2}{y^2} \right)W = 0, \quad (36)$$

with $K = -A/2$ and $\mu = \pm 1/4$. The linearly independent solutions of this equation are given by

$$W_{K,\mu}^{(1)} = e^{-y/2}y^{\mu+1/2}M(1/2 + \mu - K, 1 + 2\mu, y), \quad (37)$$

$$W_{K,\mu}^{(2)} = e^{-y/2}y^{\mu+1/2}U(1/2 + \mu - K, 1 + 2\mu, y), \quad (38)$$

where M and U are the Kummer and Tricomi functions [?], also known as confluent hypergeometric functions of the first and second kind, respectively. We will use only the function M because the function U is multivalued, and we study the solutions only for finite y since the model holds for a limited range of u values. Moreover, for $\gamma \leq 1$, the function $W_{K,\mu}$ is bounded since $M(1/2 + \mu - K, 1 + 2\mu, y)$ behaves like

$$\Gamma(1 + 2\mu) \frac{e^{y/2}y^{-1/2-\mu-K}}{\Gamma(1/2 + \mu - K)} + \frac{(-y)^{-1/2-\mu+K}}{\Gamma(1/2 + \mu + K)},$$

where Γ denotes the standard gamma function. Specifically, the Kummer function $M(a, b, y)$ is formally defined as

$$M(a, b, y) = 1 + \frac{a(1)}{b(1)}y + \dots + \frac{a(n)}{b(n)} \frac{y^n}{n!},$$

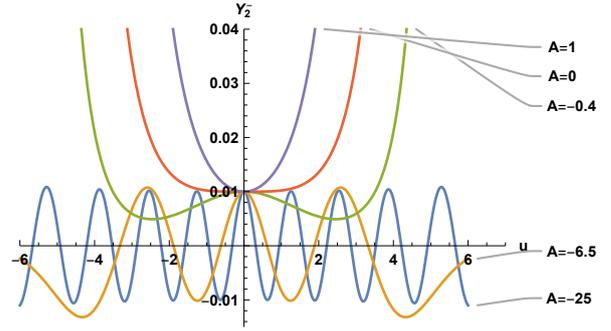


Figure 4. (Color online) Plot of the function $Y_2^-(u)$ in Eq.(40), corresponding to $\mu = -1/4$, for different values of the parameter A as in the numerical analysis of Fig.2. The integration constant c_1^- is set to match the boundary condition in $u = 0$.

where the rising factorial denoted with the index (n) is defined as

$$X_{(0)} = 1, \\ X_{(n)} = X(X + 1)(X + 2) \dots (X + n - 1).$$

The sum is converging everywhere, while the derivative of M with respect to y is given by the following recurrent equation: $dM(a, b, y)/dy = (a/b)M(a + 1, b + 1)$.

Using this formalism, we get the following solutions of Eq.(33):

$$Y_2^+(u) = c_1^+ u e^{-u^2/4} M(A/2 + 3/4, 3/2; u^2/2), \quad (39)$$

$$Y_2^-(u) = c_1^- e^{-u^2/4} M(A/2 + 1/4, 1/2; u^2/2), \quad (40)$$

where c_1^\pm are integration constants and we denoted with $+/-$ the choice $\mu = \pm 1/4$.

The zeros of the derivative of the solutions above correspond to that one of the radial component of the magnetic field from Eq.(22). Using recursive expressions, we get

$$dY_2^+/du = \\ = c_1^+ e^{-u^2/4} [(1 - u^2/2)M(A/2 + 3/4, 3/2; u^2/2) + \\ + u e^{-u^2/4} (A/3 + 1/2)u M(A/2 + 7/4, 5/2; u^2/2)], \quad (41)$$

$$dY_2^-/du = \\ = c_1^- e^{-u^2/4} (-u/2)M(A/2 + 1/4, 1/2; u^2/2) + \\ + c_1^- u e^{-u^2/4} (A + 1/2)M(A/2 + 5/4, 3/2; u^2/2). \quad (42)$$

Since in the general solution both $\mu = -1/4$ and $\mu = 1/4$ have to appear in a linear combination, we plot in Fig.4 and Fig.5 the function $Y_2(u)$ for both these values, respectively. Comparing these two pictures with the numerical integration in Fig.2, we see that the choice $\mu = -1/4$ better addresses the profile associated to the considered physical boundary conditions of the numerical integration. Moreover, it is worth underline the following behavior depending on the values of the A parameter (we focus only on the half-plane $u \geq 0$): for $-25 < A < -16$, Y_2^- oscillates; for $-16 < A < 0$, the oscillations disappear and the solution has only some maxima or minima;

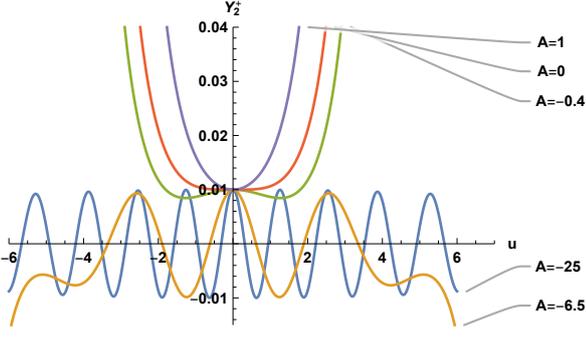


Figure 5. (Color online) Plot of the function $Y_2^+(u)$ in Eq.(39), corresponding to $\mu = 1/4$ (parameter choice as in Fig.4).

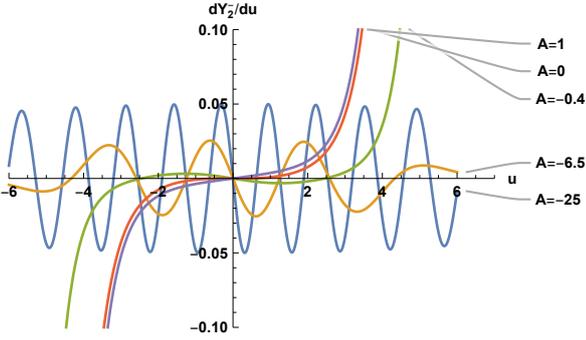


Figure 6. (Color online) Behavior of dY_2^-/du in Eq.(42) corresponding to parameter choice of Fig.4.

for $A \geq 0$ it is an increasing function. The transitions of the analytic solution are in agreement with those of the numerical solution of the previous Section. In Fig.6 we instead plot the graph of dY_2^-/du . The O -points of the radial component of the magnetic field on the vertical axis clearly emerge, placed in symmetric positions with respect to the equatorial plane. For the specific case $A = -25$, the positive zeros of B_r can be found at $u = [0, 0.6, 1.3, 1.9, 2.55, 3.2, 3.9, \dots]$. In this respect, we note that algorithms are present for systematically finding the roots of Kummer functions [?].

6.1 Asymptotic solutions

Taking into account the solutions for $Y_2(u)$ in Eqs.(39) and (40), let us now define the following functions

$$Z(A, u) = \Phi_1(u) \cos(\pi(1/4 + A/2)) + \Phi_2(u) \sin(\pi(1/4 + A/2)), \quad (43)$$

$$V(A, u) = \frac{1}{\Gamma(1/2 - A)} (\Phi_1(u) \sin(\pi(1/4 + A/2)) + \Phi_2(u) \cos(\pi(1/4 + A/2))), \quad (44)$$

where

$$\Phi_1(u) = \frac{\Gamma(1/4 - A)}{\sqrt{\pi} 2^{A/2+1/2}} Y_2^+, \quad (45)$$

$$\Phi_2(u) = \frac{\Gamma(3/4 - A/2)}{\sqrt{\pi} 2^{A/2-1/2}} Y_2^-. \quad (46)$$

In the asymptotic limit $-A \gg u^2$ (for $A < 0$) and finite u (we recall that the thin-disk hypothesis limits the range of u values), we obtain

$$Z(A, u) + i\Gamma(1/2 - A)V(A, u) = \frac{e^{i\pi(1/2+1/2A)}}{2^{(1/2+1/2A)}\sqrt{\pi}} \Gamma(1/4 - A/2) e^{\theta_r + i\theta_i} e^{ipu}, \quad (47)$$

where $p = \sqrt{-A}$, while $\theta_r(u)$ and $\theta_i(u)$ are assigned polynomials. Thus, the function $Z(A, u)$, the solution $Y_2(u)$ and its derivative with respect to u oscillate giving rise to a sequence of O -points across the vertical axis.

For positive values of A , we get instead the following asymptotic behavior for $A \gg u^2$:

$$Z(A, u) = \frac{\sqrt{\pi}}{2^{(1/2+1/2A)}\Gamma(3/4 + A/2)} e^{-\sqrt{A}u + \theta_1}, \quad (48)$$

where θ_1 is again an assigned polynomial. This case for $A > 0$ is not relevant for the present physical discussion.

6.2 Laguerre polynomials

We have shown how, in the case $-A \gg u^2$, there is a countable set of O -points, but it is also possible to find other zeros series in the correspondence of smaller $|A|$. It is easy to verify that the following relation takes place between the Kummer function $M(a, b, y)$ and the Laguerre polynomials $L_n^{(\alpha)}$:

$$M(-n, \alpha + 1, y) = L_n^{(\alpha)}(y) / \binom{n + \alpha}{n}. \quad (49)$$

For $\mu = -1/4$, Eq.(40) contains the function $M(A/2 + 1/4, 1/2; u^2/2)$, and we easily get the specific values: $n = -A/2 - 1/4$ and $\alpha = -1/2$. We thus obtain the following general expression for the derivative of Y_2^- :

$$dY_2^-/du = -c_1^- e^{-u^2/4} \times \left(u L_{n-1}^{(1+\alpha)}(u^2/4) + (u/2) L_n^{(\alpha)}(u^2/4) \right) / \binom{n + \alpha}{n}. \quad (50)$$

The equation above is a particular form of Eq.(42) and, since it is well-known that for each n there is a finite number of zeros of the Laguerre polynomials which increases with n , we have shown how a countable set of zeros on the z axis emerges. In particular, the zeros of Laguerre polynomial $L_n^{(\alpha)}$ belong to the interval $(0, n + \alpha + (n-1)\sqrt{n + \alpha})$, and in Fig.7 we plot, as an example, dY_2^-/du from Eq.(50) for the first 10 integer n values.

7 Discussion

We have shown how the z -dependence of the magnetic field is associated to O -points of the corresponding configuration. In such points, the radial component of the magnetic field, due to the plasma backreaction only, vanishes.

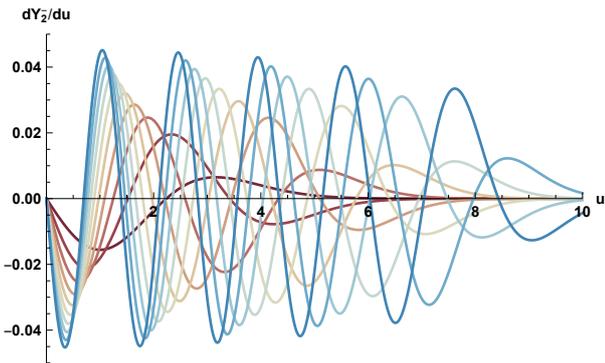


Figure 7. (Color online) Plot of dY_2^-/du from Eq.(50) (zoom of the half-plane $u > 0$) by setting integers n in $[1, 10]$. Standard color scheme from red ($n = 1$) to blue ($n = 10$). The number of oscillations increases with increasing n .

Here, we outline a new feature induced by the crystalline morphology of the plasma response to the dipole field of the central object.

Let us now discuss the relevance of the O -points derived above, in view of the physical properties of the plasma disk thought as an accreting structure. Our analysis has been performed for an ideal plasma disk in the absence of poloidal components of the velocity field, *i.e.*, $v_r = v_z \equiv 0$. However, in the Standard Model of accretion [?], the plasma must have a finite electric conductivity and the azimuthal component of the generalized Ohm law reads

$$v_z B_r - v_r B_z = J_\phi / \sigma, \quad (51)$$

where J_ϕ is the corresponding azimuthal component of the current density and σ is the constant coefficient of electric conductivity. In the standard mechanism for accretion onto a compact object, both the quantities B_r and v_z are essentially negligible and, where σ takes very large values, the accretion is suppressed, *i.e.*, since $B_z \neq 0$, we must get $v_r \simeq 0$.

In the spirit of the present analysis, in those regions where the plasma is quasi-ideal, we find the relation

$$v_z = v_r B_z / B_r. \quad (52)$$

Clearly, nearby the numerous O -points of the magnetic profile, where B_r is nearly vanishing, we get very high values of the vertical velocity, which suggests the existence of privileged plasma sites for the jet formation. This perspective has been investigated in some detail in Refs.[?,?].

The main merit of the present analysis consists of having outlined the general character that the O -points takes in the linear profile of the perturbed magnetic field, as soon as the short wavelength backreaction of the plasma is excited in disk with high β values.

8 Conclusions

We analyzed general features of the small scale morphology of the backreaction that is generated in a thin plasma,

embedded in the magnetic field of the gravitationally confining central object. In particular, we studied the Master Equation of the so-called crystalline structure [?], searching for a satisfactory characterization of the admissible profile in the linear backreaction limit.

We have followed two different, but complementary, approaches. The first one is based on a general Fourier expansion, while the second one relies on a separated variable procedure. This analysis offered a complete spectrum of the available solutions, outlining new behaviors in the plasma equilibrium, absent in the basic treatment in Refs.[?,?].

The main merit of this systematic study of the linear backreaction of the plasma disk consists of the relevant physical implications of the obtained solutions, as discussed above, *i.e.*, the inverse behavior of the vertical gradient with respect to the background one (as the highest harmonics are considered) and the possible emergence of a consistent sequence of O -points, where the radial magnetic field component vanishes.

The first of these issues is of interest in view of the non-linear scenario of the plasma backreaction discussed in Refs.[?,?], where, see also Ref.[?], the possibility for the radial fragmentation of the disk into a microscopic array of rings is demonstrated. In this respect, the linear inversion of the vertical gradients suggests that, in the non-linear regime, a separation of the disk into two symmetric components, up and down the equatorial plane, could take place, with significant implications on the transport features across the disk.

The second result, concerning the appearance of O -point series of the magnetic configuration, is relevant in view of the realization of conditions for jet emission. In fact, nearby the O -points, when the plasma resistivity is suppressed, the vertical component of the plasma velocity can take very large values, see the discussion in Refs.[?,?] about the seeds of a vertical matter flux from the disk. Using the solution given in terms of hypergeometric confluent functions, we have also analytically shown the existence of the O -point series and such a series critically depends on the parameter of the separation of variable technique.

The systematic analysis here pursued of the Master Equation for the crystalline structure of a stellar accretion disk, constitutes a well-grounded starting point for the investigation of non-linear features of the magnetic confinement of the plasma disk. In fact, as discussed in Ref.[?], the extreme non-linear regime is characterized by a dominant influence of the MHD-force in confining the plasma with respect to the radial gravitational field. Furthermore, in Ref.[?] it is outlined how the request that the Master Equation holds in the non-linear case too turns out to be a proper choice to close the equilibrium system, preserving the crystalline morphology.

9 Authors contributions

All the authors were involved in the preparation of the manuscript. All the authors have read and approved the final manuscript.